

# LABORATORY ASTROPHYSICS WHITE PAPER

## SUMMARY OF LABORATORY ASTROPHYSICS NEEDS (DRAFT June 12, 2002)

### BASED ON DISCUSSIONS HELD AT THE NASA 2002 LABORATORY ASTROPHYSICS WORKSHOP (NASA Ames Research Center, May 1-3, 2002)

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The NASA Laboratory Astrophysics Workshop (NASA LAW) met at NASA Ames Research Center from 1-3 May 2002 to assess the role that laboratory astrophysics plays in the optimization of NASA missions, both at the science conception level and at the science return level. Space missions provide understanding of fundamental questions regarding the origin and evolution of galaxies, stars, and planetary systems. In all of these areas the interpretation of results from NASA's space missions relies crucially upon data obtained from the laboratory.

We stress that Laboratory Astrophysics is important not only in the interpretation of data, but also in the design and planning of future missions. We recognize a symbiosis between missions to explore the universe and the underlying basic data needed to interpret the data from those missions.

In the following we provide a summary of the consensus results from our Workshop, starting with general programmatic findings and followed by a list of more specific scientific areas that need attention. We stress that this is a "living document" and that these lists are subject to change as new missions or new areas of research rise to the fore.

## **I. General findings:**

- The number and scope of NASA missions requiring supporting laboratory data (especially spectroscopy) has risen dramatically, but the funding for Laboratory Astrophysics has remained flat. The current funding profile cannot meet the needs of existing missions or prepare for future missions.
- Critical compilation of databases was identified as a high priority need. For example, NIST has been maintaining the highest quality atomic database in the world, a resource relied upon for basic data in several communities, but for which NIST does not provide internal support. The over-strained NASA budget has not been able to provide the level of support required to maintain this database, and resources from other agencies have also diminished. We are concerned that such databases may not survive.
- Laboratory facilities are aging and major funding is required to replace them with modern, state-of-the-art equipment. Unless new facilities can be built it will become increasingly difficult for laboratory astrophysics to keep up with the demand for atomic and molecular data to support space-based and ground-based observatories.
- The training of new scientists in Laboratory Astrophysics is crucial for the future of the field, but the low level of funding is making it increasingly more difficult to attract students and train the scientists of tomorrow to provide the required laboratory data. Retention of existing groups is also a concern as it is dependent on continuing sources of grant funds.
- Within the next decade four major infrared and/or submillimeter astronomical observatories will be launched, whose success will depend largely on accurate and comprehensive laboratory data on atomic and molecular transitions in these spectral regions, which are to date largely unexplored in the lab. Lab data in these spectral regimes are crucial to the success of these missions.
- Major needs exist for laboratory data to help in the interpretation of spectra from current and planned UV and X-ray missions. This need extends also to upcoming IR and submillimeter missions, which will observe highly-redshifted UV features.
- Spectroscopy is of the essence. Until the 1990s, sensitivities in most spectral ranges were too low to permit high resolution spectroscopy, and missions like IRAS and ROSAT concentrated on broad-band mapping and photometry. The situation is drastically different today as recent, current, and planned future missions provide high-resolution spectra to look at detailed chemical and physical processes on solar system, stellar, galactic, and extragalactic scales over a wide range of photon energies.
- The demands on laboratory astrophysics for high signal-to-noise and high spectral resolution are unprecedented. Many of NASA's planned missions have spectroscopic capabilities with exquisite sensitivity, and most will have extremely high spectral resolving power. A host of new spectral features numbering in the millions will need to be identified and interpreted.

To meet the objectives of these planned missions and to make their findings accessible to the community in clear, understandable terms, we will need the support of a coherent and vigorous Laboratory Astrophysics program that integrates theory, modeling, and experiment.

## II. Needs for laboratory data supporting specific missions and spectral bands.

We now consider the more specific scientific connections between laboratory astrophysics and various current and planned NASA missions. These needs for laboratory data differ significantly among wavelength ranges. Hence the various spectral regions are considered here in order of increasing wavelength. Within each wavelength interval we specify the current or planned missions whose data interpretation depends on the results of Laboratory Astrophysics studies.

### *The X-Ray Spectral Region*

*Missions:* Chandra, XMM-Newton, Astro-E2, and Constellation-X

X-ray observations can be used to address a number of fundamental questions in astrophysics. Relativistically broadened metal lines in AGN and quasars can be used to study black holes. High resolution spectra of AGN, quasars, XRBs, and CVs help us learn about accretion in the vicinity of compact objects. X-ray emission from infrared bright galaxies tell us about the relationship between the AGN and starburst components of galaxies. Stellar winds, supernova remnants, and the ISM can be used to study nucleosynthesis and the evolution of our Galaxy and the universe. Observations of galaxy clusters and the IGM provide information on the formation of large scale structure in the universe, the formation of galaxies, and provide constraints on the dark matter component of the universe. Stellar coronal observations can be used to study the connection between the stellar photosphere and corona and the physical processes involved in heating and mass supply to stellar coronae.

X-rays provide a quantitative understanding of highly ionized plasmas, both electron-ionized and photoionized. Electron-ionized plasmas are formed in stellar coronae, supernova remnants, the ISM, galaxies, and clusters of galaxies. Photoionized plasmas are found in planetary nebulae, H II regions, the IGM, AGN, XRBs, and CVs. Interpreting X-ray spectra from all these sources depends critically upon input from laboratory astrophysics. Priorities for measurements include atomic rate coefficients for low and high temperature dielectronic recombination, state-specific and total charge transfer for multiply charged ions on H and H<sub>2</sub>, collisional excitation, inner shell photoabsorption, and photoexcitation and photoionization cross sections as well as fluorescence and Auger yields. Measurements of fundamental spectroscopic quantities are also needed, including M- and L-shell emission and satellite line identifications and wavelengths (accurate to one part in 10<sup>5</sup>) for Ne, Mg, Si, S, Ar, Ca, Fe, and Ni.

While most atomic data used to interpret X-ray plasmas are compiled from theoretical calculations, laboratory measurements are needed to determine the accuracy of rates for critical diagnostics. As an example, the strong set of Fe XVII X-ray emission lines provides information on temperatures, densities, and opacities, but, despite recent laboratory studies, their astrophysical interpretations remain unclear. Laboratory surveys of the X-ray spectral content are also needed to provide critical tests of the completeness of spectral models. For moderate resolution X-ray data, complete models are required for global spectral fitting; at higher resolution, completeness studies are necessary in order to assess blending around diagnostic lines.

The atomic data for collisionally ionized plasmas, relevant to stellar coronae, galaxies, and clusters of galaxies, have received the best critical evaluations and laboratory tests. Non-equilibrium effects of X-ray spectra, such as those expected in supernova remnants, require further studies, as the atomic rate coefficients are often less accurate away from equilibrium temperatures. Priorities for future work need to include critical tests of the atomic data used for X-ray photoionized plasmas, such as found in accretion disks and black hole environments. Since spectral models for photoionized sources need to make assumptions about the astrophysics (e.g., energy balance and geometry) to incorporate atomic physics, it is critical that laboratory studies benchmark the essential atomic data, therefore studies of low temperature dielectronic recombination should receive high priority.

New results from Chandra and XMM-Newton suggest additional areas in need of laboratory astrophysics. Recent astrophysical observations have tentatively identified X-ray absorption by molecules, a new area for laboratory measurements of edge physics, which can lead to differentiation between gas and dust in diffuse media. Closer to home, objects in the solar system, such as comets, the Jovian aurora, the Io plasma torus, and the Jovian Galilean satellites (Io, Europa, Ganymede), emit X-rays. Cometary X-ray emission is tentatively attributed to charge transfer between the solar wind ions and the comet. If charge transfer is the correct mechanism, these observations provide a sample of the comet coma material as well as information on the solar wind velocity and elemental charge states. Accurate cross sections for charge transfer are needed, to compute both state-specific partial rates and total rates, for multiply charged ions on H and H<sub>2</sub>. Laboratory measurements of charge exchange cross sections will help to validate the models for solar system X-ray emission, and may be of importance in X-ray photoionized plasmas as well.

*Summary of needs for lab data to support X-ray astrophysics:*

- Atomic rate coefficients for:
  - + low and high temperature dielectronic recombination;
  - + state-specific and total charge transfer for multiply charged ions on H and H<sub>2</sub>;
  - + collisional excitation and ionization;
  - + inner shell photoabsorption;
  - + photoexcitation and photoionization cross sections;
  - + fluorescence and Auger yields.
- Measurements of fundamental spectroscopic quantities:
  - + M- and L-shell emission and satellite line identifications and wavelengths (accurate to one part in 10<sup>5</sup>) for Ne, Mg, Si, S, Ar, Ca, Fe, and Ni.
- Measurements of X-ray absorption spectra due to molecules.

### ***The UV Spectral Region***

***Missions:*** HST(STIS), HST(COS), and FUSE, for low redshift; SOFIA, Herschel, and NGST for higher redshifts

This spectral region will provide ages and metallicities of old galaxies, the sequence of galaxy formation, tests of cosmological models, and an understanding of the relative importance of r- versus s-process nucleosynthesis of neutron capture elements as a function of time since the Big Bang. The UV region also provides insights on chemically/isotopically peculiar stars, on segregation of elements in stellar photospheres, and on mass transfer in binary systems. These problems need comprehensive classified spectral line lists in the UV, including wavelengths accurate to 1 part in  $10^7$ , transition probabilities accurate to 5 to 10%, and hyperfine/isotopic parameters for high abundance Fe group elements in the first three stages of ionization, and similar spectral line lists in the UV for strong lines of lower abundance elements including the neutron capture elements.

We draw special attention to the far-UV spectral region where FUSE operates. Here special coatings and detectors are needed, with the result that lab data are lagging far behind the astronomical data. The lack of experimental data in this spectral region has hampered progress in theoretical studies as well as the interpretation of astronomical data.

The UV region is key to interpreting visible and UV spectra of important interstellar molecules, such as large organic species that carry the ubiquitous IR emission bands (UIBs) and diffuse interstellar bands (DIBs) and that may be related to the origin of life. Studies of the UV characteristics of such molecules and their dependence on molecular structure and charge state is of key importance for our understanding of this ubiquitous molecular component of the ISM. Identification of UV spectra of large aromatics, and in particular of PAHs, is especially important to address these issues and represents one of the key science goals of the HST(COS). UV spectra are uniquely capable of identifying specific molecules, in contrast with the less specific transitions observed in the IR. Lab studies provide spectroscopy of large organic molecules (such as PAHs) and their ions in the solid and in the gas phases, measurements of chemical reaction rate coefficients, and recombination cross sections. This work must be complemented by quantum theory calculations so that the lab data are properly interpreted.

The molecule CO is a key component of dense interstellar clouds and a probe of local interstellar conditions. Measurements are needed of oscillator strengths for UV intersystem bands for CO and its isotopomers, improved wavelength measurements in the UV, and photodissociation cross sections including their J-dependence at appropriate interstellar temperatures.

Astronomers also require an improved understanding of the energetics of interstellar dust in a variety of environments (photon-pumping mechanisms, etc.). Such insight arises from spectroscopic signatures that provide direct information on the composition and evolution of dust. Previous studies in the UV have focused on the only identified spectral feature (at 2200 Å), but all materials should show UV spectral signatures. Laboratory measurements of the optical properties of bulk materials such as carbonaceous solids, metallic carbides, sulfides, oxides, as well as weak features of other common materials, are needed, as well as studies of the properties of very small (nano-sized) particles of the same substances, which differ from the bulk properties.

The UV wavelength region, often used in conjunction with other wavelengths, provides an understanding of the fundamental processes (and especially the energy balance) associated with emissions from planetary atmospheres and magnetospheres, including planetary aurora and dayglow emissions (relevant for all planets and satellites with atmospheres and magnetospheres), as well as comets. Laboratory studies must provide reaction rates and electron impact excitation rate coefficients.

Light reflected from atmospheres of planets, in which absorption from the planetary atmosphere is superposed upon the solar spectrum, leads to a determination of the composition and structure of planetary atmospheres and interiors. Different wavelengths probe various depths in the atmosphere. Laboratory studies must provide line lists for the parent species (primarily those arising from methane, water, and carbon monoxide), extensive sets of rate coefficients for all classes of reactions used in the photochemical models, and an understanding of the radiative transfer and temperature effects on line shapes. Low temperature rate constants are needed in many cases. It is not known how the products of many photodissociation rates are distributed in energy states. Equations of state, solubility, and molecular diffusion in H<sub>2</sub>/He mixtures at low temperature and high density are needed for studies of giant planets that can also be used as a basis for understanding brown dwarfs.

Lack of reflectance spectra (UV-visible-NIR) of low temperature frosts/volatile ices has inhibited interpretation of the Galileo data. Unless something is done in the near future, the situation will be similar for Saturn Cassini data. Water is reasonably well covered, and the mid- and far-IR has been done for astrophysical ices, although not at the 50-150K temperatures relevant for solar system objects. Optical constants/properties of organic solids (important for most "red" solid bodies in the outer solar system) and of solid sulfur are needed.

*Summary of needs for lab data to support UV astrophysics:*

- Comprehensive line lists and f-values for atoms and ions in all stages, for: rare elements seen in stars (e.g., neutron-capture elements);
  - + ionization stages not previously studied;
  - + the far-UV spectral region.
- Molecular spectra of small molecules, including:
  - + CO (and CO isotopomer) intersystem bands;
  - + large organics such as PAHs and heterocycles;
  - + molecules expected to be abundant in planetary atmospheres;
  - + other species possibly detectable at high S/N in astronomical spectra.
- Spectra and ionization analyses of large organic molecules such as PAHs and derivatives.
- Photodissociation cross sections of important interstellar molecules such as CO and H<sub>2</sub>O.
- UV pumping rates leading to IR emission by molecules.
- Optical, physical, and chemical and properties of dust analogues, including:
  - + optical properties of bulk materials such as carbonaceous solids, metallic carbides, sulfides, oxides;
  - + weak spectral features of other common materials;

- + comparative physical, chemical, and spectroscopic properties of interstellar and interplanetary particles
- + formation processes and rates for molecule formation (especially molecular hydrogen) on grain surfaces;
- + studies of nano-sized particles.
- Spectra and excitation cross-sections of molecules important in planetary atmospheres, particularly for species arising from methane, water, and carbon monoxide; and sulfur-bearing species.
- Reflectance spectra (UV-visible-NIR) of low temperature frosts/volatile ices relevant for studies of satellites of the outer planets.
- UV-visible optical constants/properties of organic solids (important for most "red" solid bodies in the outer solar system) and of solid sulfur.

### ***The Infrared and Sub-millimeter Spectral Region***

***Missions:*** HST(NICMOS), SIRTF, SOFIA, Herschel, and NGST

The infrared and submillimeter wavelength region is of paramount importance for studies of the cold, molecular and dusty universe. The formation of stars and planetary systems takes place deep inside cold gas and dust clouds, often obscured by hundreds of visual magnitudes of extinction. At high redshifts, the assembly of galaxies through the merging of smaller units is accompanied by large amounts of obscuring dust. In order to penetrate these dusty regions and probe the processes occurring deep within, observations at infrared and submillimeter wavelengths are essential. Moreover, these wavelengths provide sensitive probes of the physical conditions in and dynamics of such regions through the pure rotation and ro-vibrational transitions of small molecules as well as the detailed profiles of transitions in larger molecules and those associated with dust. NASA will launch and participate in a number of missions centered on this wavelength region (SIRTF, SOFIA, Herschel, NGST) which will chart the star formation history of the universe, star and planet formation in the Milky Way, the galactic life cycle of the elements, and the molecular universe, in general. Together these span most of the key questions in modern astronomy.

The ensemble of current and planned IR/sub-mm missions will bring in enormous quantities of data in spectral regions where little is known. Laboratory studies are essential in order to support the analyses of these data. Accurate frequency measurements and band analyses are needed for atomic and molecular species having transitions between 500 GHz and 2 THz, a largely unexplored spectral region important for many kinds of astrophysical studies, including low- and high-density interstellar gases and plasmas, circumstellar clouds and disks, cool star atmospheres, and star- and planet-forming regions. Studies of interstellar dust will also be especially important.

Mid-IR spectra of individual objects such HII regions, reflection nebulae, and planetary nebulae as well as the general interstellar medium of galaxies as a whole are dominated by a set of mid-IR emission features due to large PAH molecules. Studies of the IR characteristics of such molecules and their dependence on molecular structure and charge state is of key importance for our understanding of this ubiquitous molecular component of the ISM.

At long wavelengths, the continuum dust opacity is uncertain by an order of magnitude. IR/sub-mm transitions of interstellar dust grains must be used to determine their specific mineral composition, hence their opacities, which determine inferred grain temperatures and the masses of dusty objects, including the interstellar medium of entire galaxies. Emission bands from warm astronomical environments such as circumstellar regions, planetary nebulae, and star-forming clouds lead to the determination of the composition and physical conditions in regions where stars and planets form. The compositions of cool stars provide the nucleosynthetic history of our galaxy and of others.

The laboratory data essential for investigations of dust include measurements of emission and absorption spectra (opacities at various wavelengths, especially in the FIR) of candidate grain materials (both carbonaceous and siliceous). Metallic carbides, sulfides, oxides, as well as weak features of common materials, are important. For abundant materials (e.g., forms of carbon such as PAHs), the measurements should range from molecules to nano-particles to bulk materials.

Also vital are measurements of molecular wavelengths and transition probabilities in the near-IR, where cool stars emit most of their flux. There are many as-yet unexploited transitions of important species.

*Summary of needs for lab data to support IR and sub-mm astrophysics:*

- Frequency measurements and band analyses for molecular species whose transitions lie between 500 GHz and 2 THz.
- Accurate wavelengths and transition probabilities of far-IR/sub-mm atomic fine structure transitions that are key diagnostics of conditions in low-density environments.
- Spectra and transition probabilities for small molecules such as light hydrides and metal hydrides which are important indicators of composition and which also dominate the cooling in low-density environments.
- Spectra and ionization analyses of large organic molecules such as PAHs.
- IR/sub-mm spectra of solid-state transitions in candidate grain materials:
  - + for metallic carbides, sulfides, oxides, as well as weak features of common materials;
  - + specific mineral compositions;
  - + dust opacities;
  - + optical properties for sizes ranging from molecules to nano-particles to bulk materials.
- Near-IR spectra and line strengths for atomic species and isotopes needed for abundance analyses in cool stars, for common and rare elements and isotopes.
- Reflectance spectra (UV-visible-NIR) of low temperature frosts/volatile ices relevant for studies of satellites of the outer planets.

***Non-Wavelength-Specific needs involve fundamental physical and chemical processes that affect the scientific output of all NASA missions***

Understanding of celestial objects requires knowledge of the rates of relevant physical and chemical processes. Only laboratory data can provide reaction rates, sticking coefficients, and desorption rates for processes in the gas phase, on the surfaces of small (nano-sized) and larger (micro-sized) grains, within interstellar ices, and on solar system objects such as icy satellites and Kuiper Belt objects. For instance, a quantitative understanding of the formation of the most important interstellar molecule, H<sub>2</sub>, is very important and still incomplete. The quantitative understanding of the energy balance of the ISM (from the diffuse component to dense clouds) and the resulting phase structure requires studies of photoelectric yield and recombination cross sections of astrophysically relevant molecules and dust materials. Energetic processing of molecules, ices and dust materials by particle bombardment, and FUV and X-ray irradiation can substantially modify the composition and structure of these materials.

Low velocity collisions of icy dirt balls and mechanical properties of cryogenic porous ice/rock aggregates have direct relevance to the Kuiper Belt. The limitations on the available lab data will also hamper interpretations of thermal IR spectra to be obtained from the surface in 2004 by the Mars Exploration Rovers. The situation is similar in the visible to NIR, which will be studied extensively at the tens-of-meters scale by a NIR spectrometer on the 2005 Mars Reconnaissance Orbiter mission.

Web-accessible, critically evaluated databases should become available. The efficiency gains obtained by eliminating the need for individual scientists to search original literature for data are significant. Critical evaluation and the establishment of reliable error bars on the data are important. Reliable error bars add much confidence to scientific conclusions based on the interpretation of astrophysical observations with the data and modeling with the data.

***Summary of needs for lab measurements of physical and chemical processes:***

- Atomic rates and cross sections, especially for processes relevant to X ray properties of plasmas (see above).
- Quantitative studies of processes relevant to the formation of molecular hydrogen and other molecules on grains.
- Chemical reaction rates for large and small molecules of interest for all astrophysical environments.
- Charge balance measurements (ionization and electron recombination cross sections), especially for large organic molecules such as PAHs.
- Photoelectric yield and recombination cross sections of astrophysically relevant dust materials.
- Processing of molecules, ices, and dust materials by particle bombardment and FUV and X-ray irradiation.
- Low velocity collisions of icy dirt balls and mechanical properties of cryogenic porous ice/rock aggregates.
- Development and maintenance of web-accessible, critically evaluated data bases.